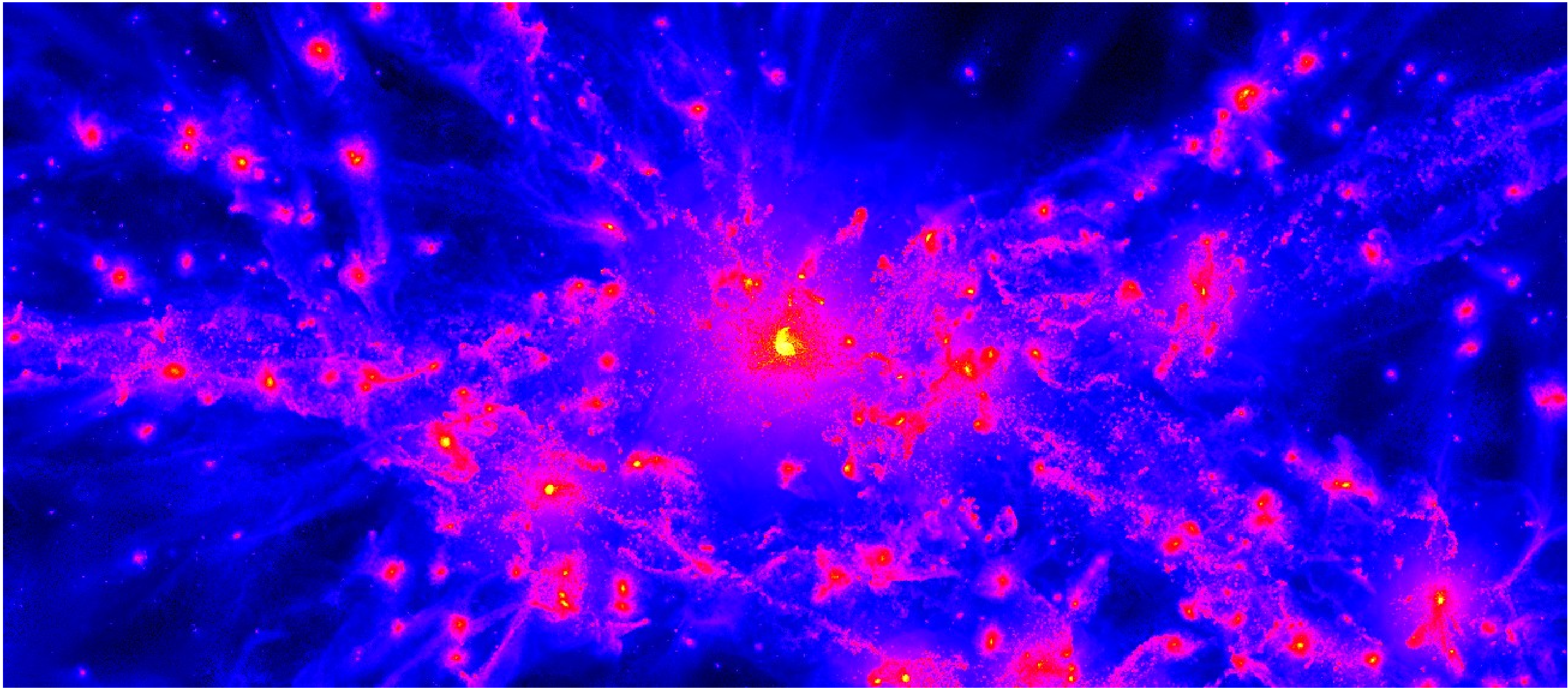
Unified Modeling of Galaxy Populations in Clusters



Thomas Quinn
University of Washington
NSF PRAC Award 1613674

Fabio Governato
Lauren Anderson
Michael Tremmel
Ferah Munshi
Joachim Stadel
James Wadsley
Greg Stinson

PPL PARALLEL PROGRAMMING LAB
DEPT OF COMPUTER SCIENCE, UNIVERSITY OF ILLINOIS

Laxmikant Kale
Filippo Gioachin
Pritish Jetley
Celso Mendes
Amit Sharma
Lukasz Wesolowski
Gengbin Zheng
Edgar Solomonik
Harshitha Menon
Orion Lawlor

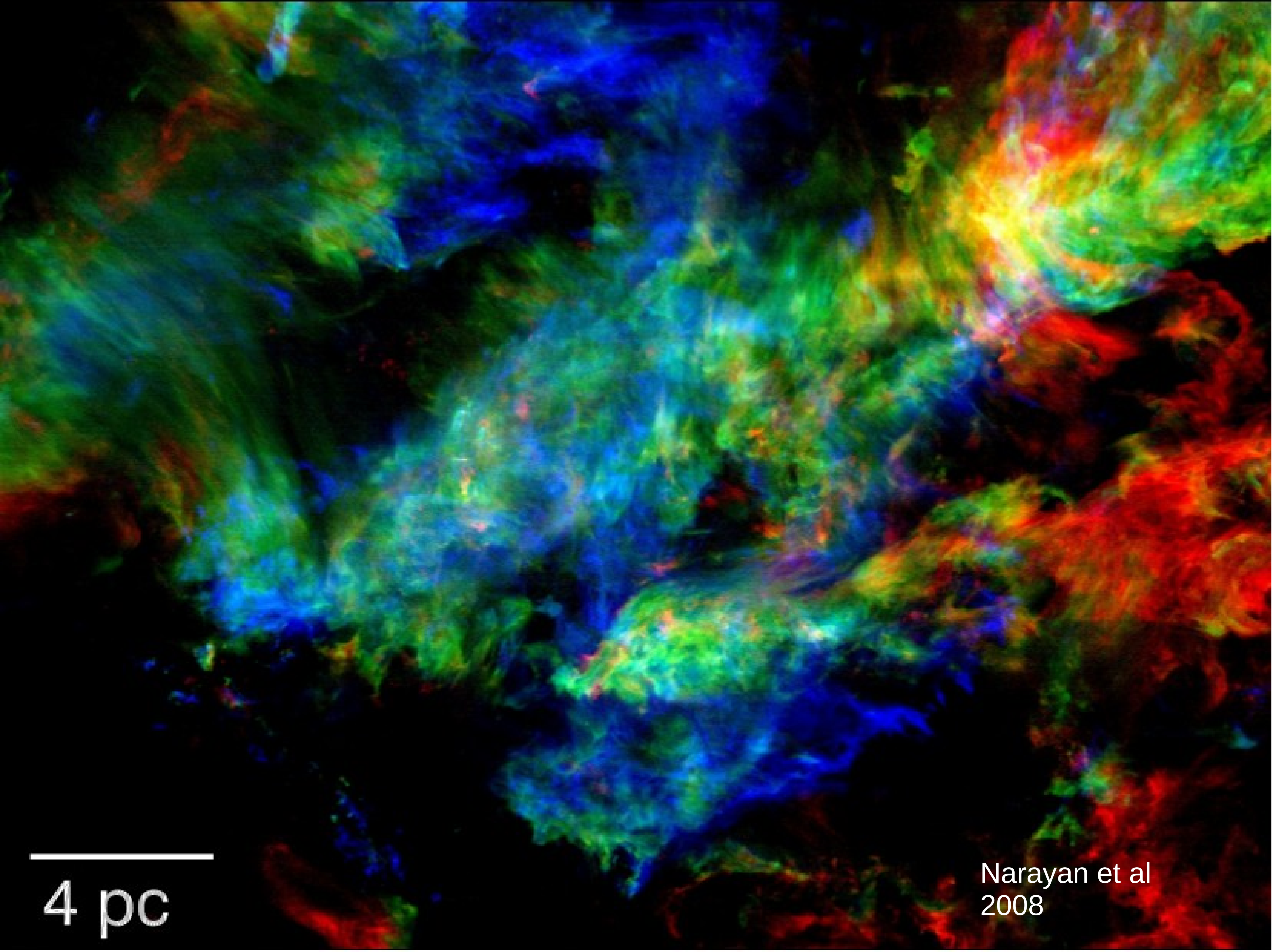
Outline

- Scientific background (Why it matters)
- Need for high resolution (Why Blue Waters)
- Previous Results (Accomplishments)
- The Cluster Clustering Problem (Key Challenges)
- Charm++ and ChaNGa (Key Challenges)
- Preliminary results (Accomplishments)

Galaxies: can we form one of these?

Modeling Star Formation: it's hard

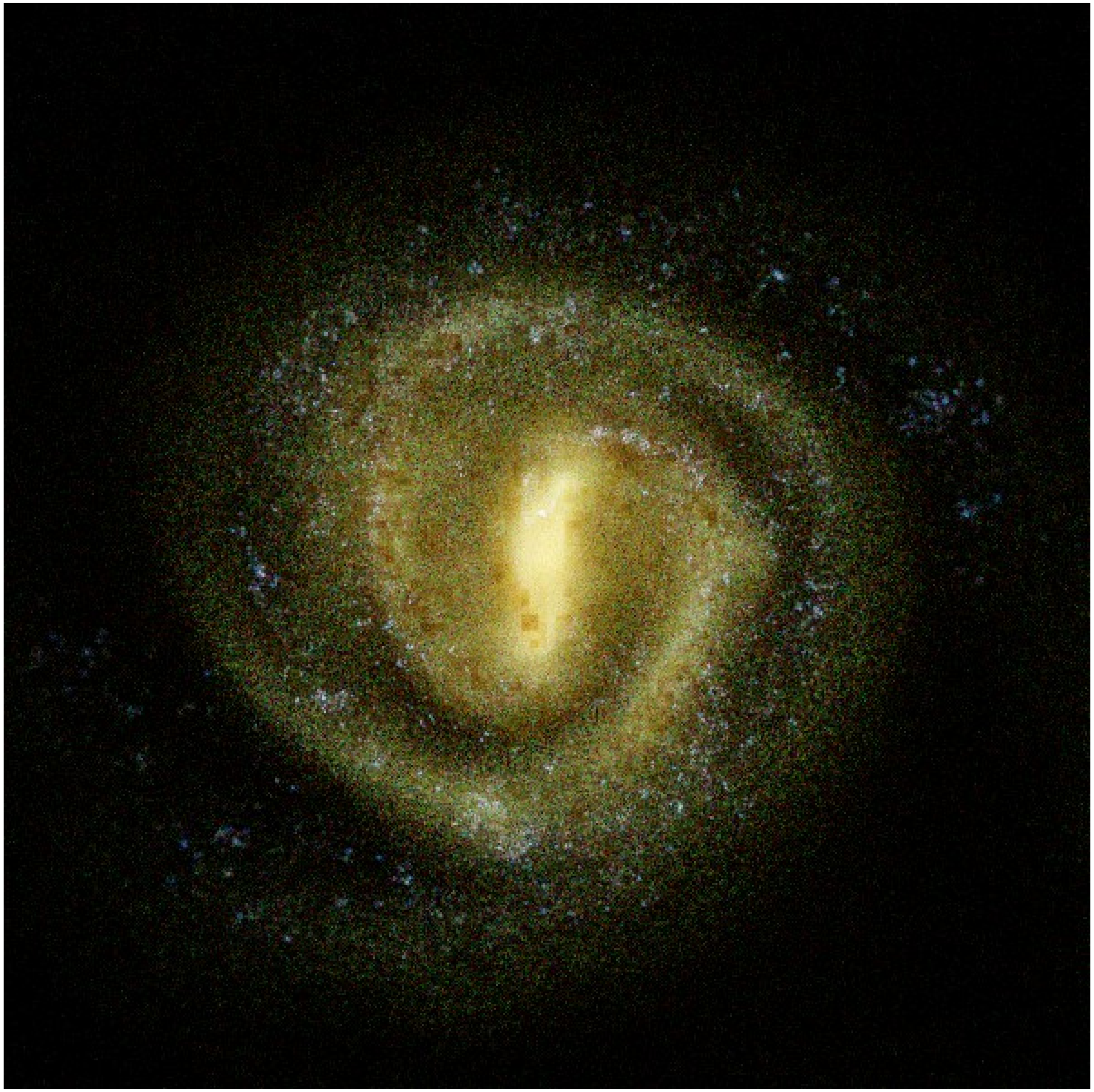
- Gravitational Instabilities
- Magnetic Fields
- Radiative Transfer
- Molecular/Dust Chemistry
- Driven at large scales: differential rotation
- Driven at small scales: Supernovea and Stellar Winds
- Scales unresolvable in cosmological simulations


4 pc
Narayan et al
2008

Resolution and Subgrid Models

- Maximize Simulation Resolution
  - Capture tidal torques/accretion history (20+ Mpc)
  - Adapt resolution to galaxy (sub-Kpc, 10^5 Msun)
- Capture Star Formation in a sub-grid model
  - Stars form in high density environments
  - Supernovea/stellar winds/radiation regulate star formation
  - Mitigate issues with poor resolution (overcooling)
  - Tune to match present day stellar populations

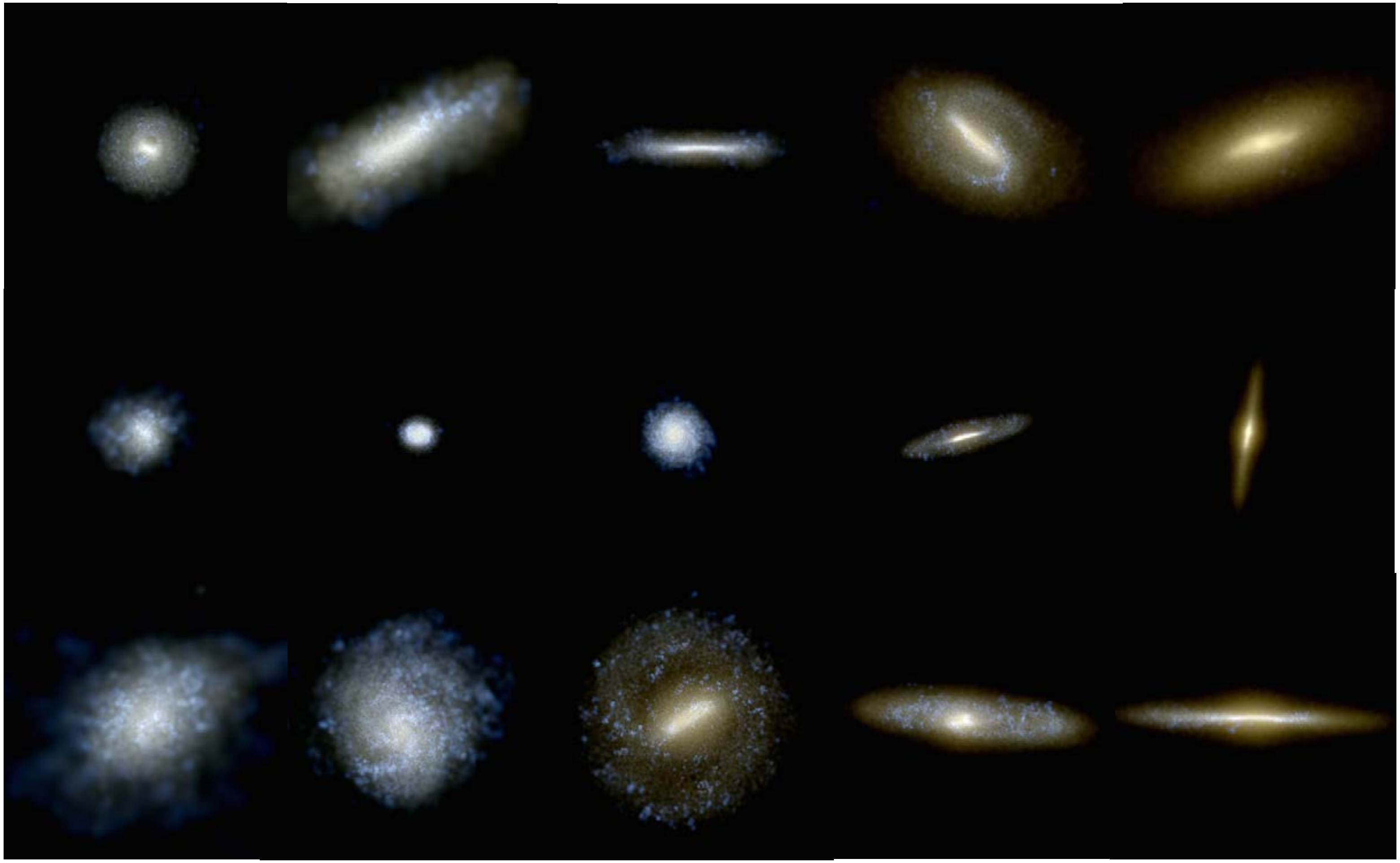
Previous PRAC: good morphologies



Danielle
Skinner

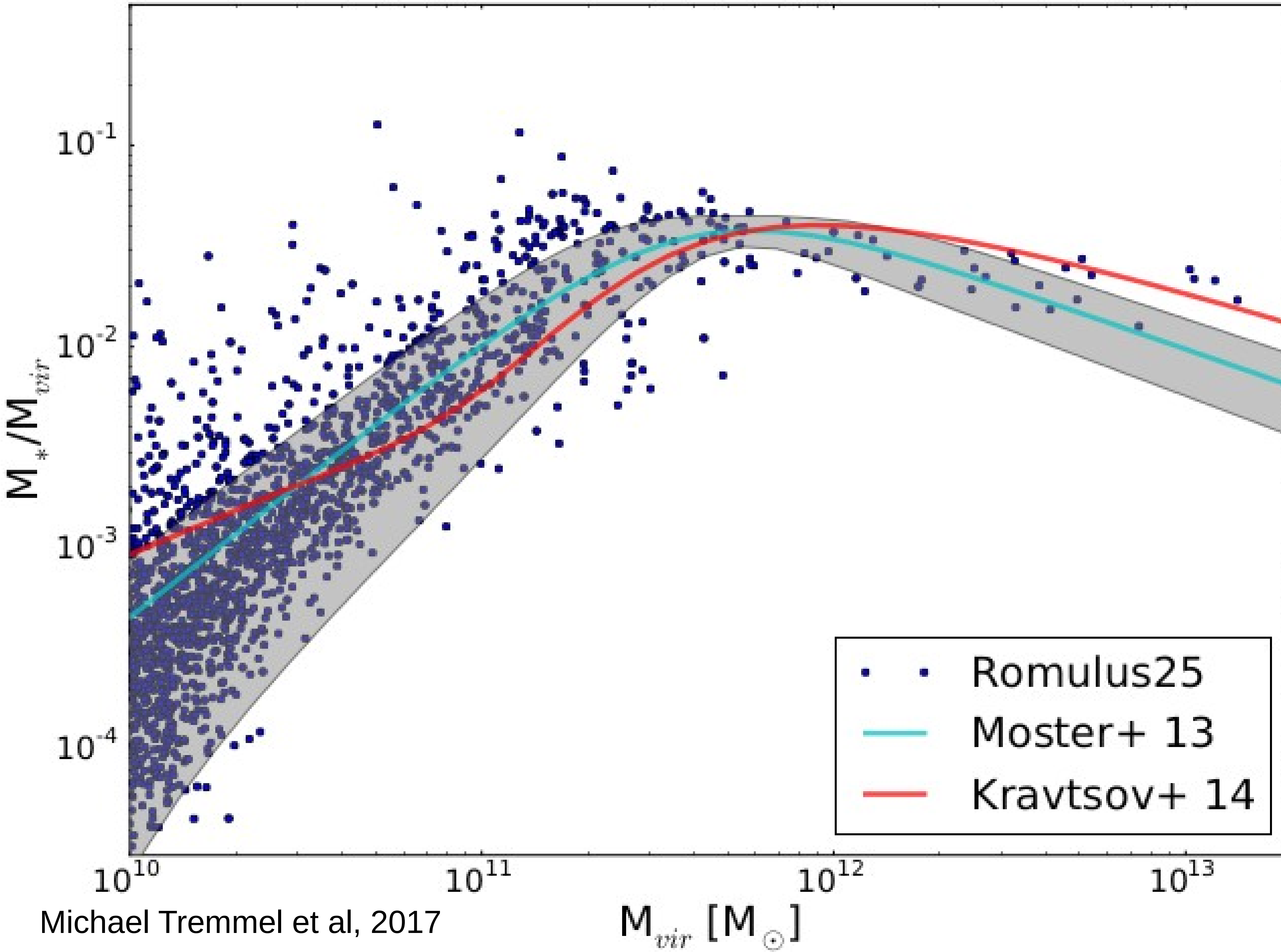
Good morphologies across a population

z = 3 z = 2 z = 1.2 z = 0.75 z = 0.5

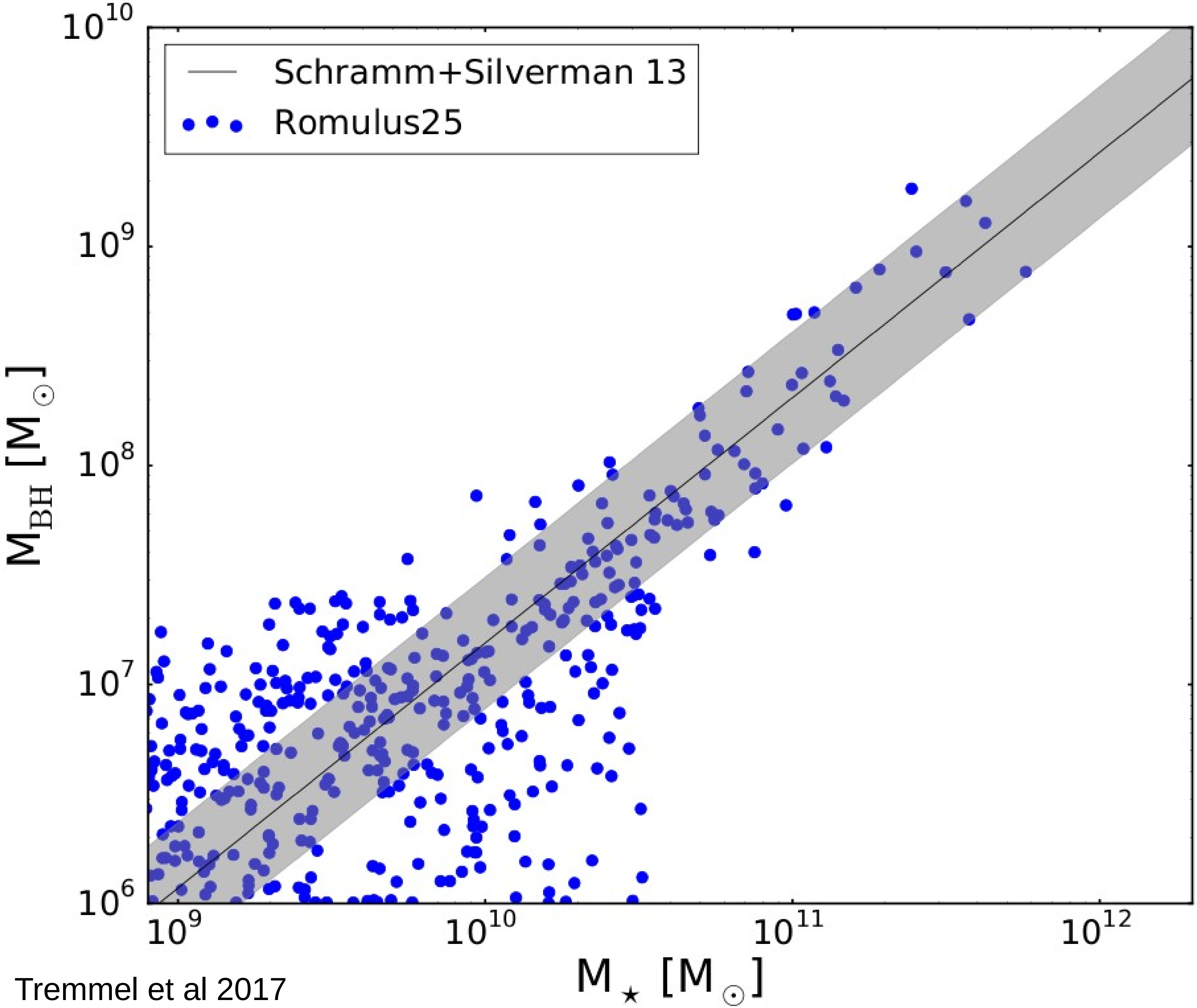


Black hole/AGN feedback

- Supernova feedback doesn't suppress star formation in massive galaxies
  - Modeling of more energetic feedback required
- Components of AGN modeling:
  - Seed (1e6 Msun) BH form in dense, low metallicity gas
  - BH grow from accreting gas, and release energy into the surrounding gas (Active Galactic Nuclei)
  - BH in merging galaxies sink to the center and merge (LIGO, eLISA)

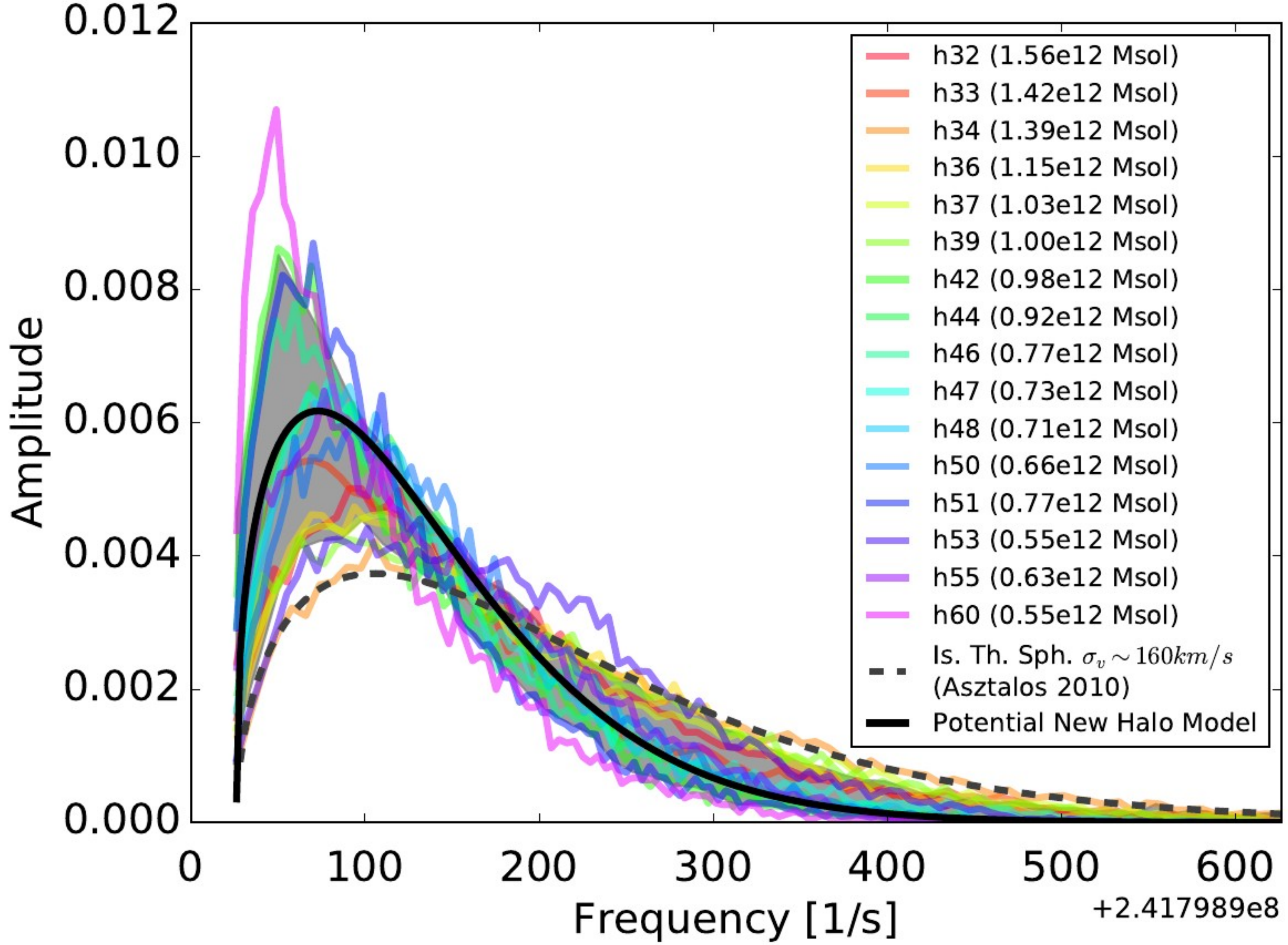


Michael Tremmel et al, 2017



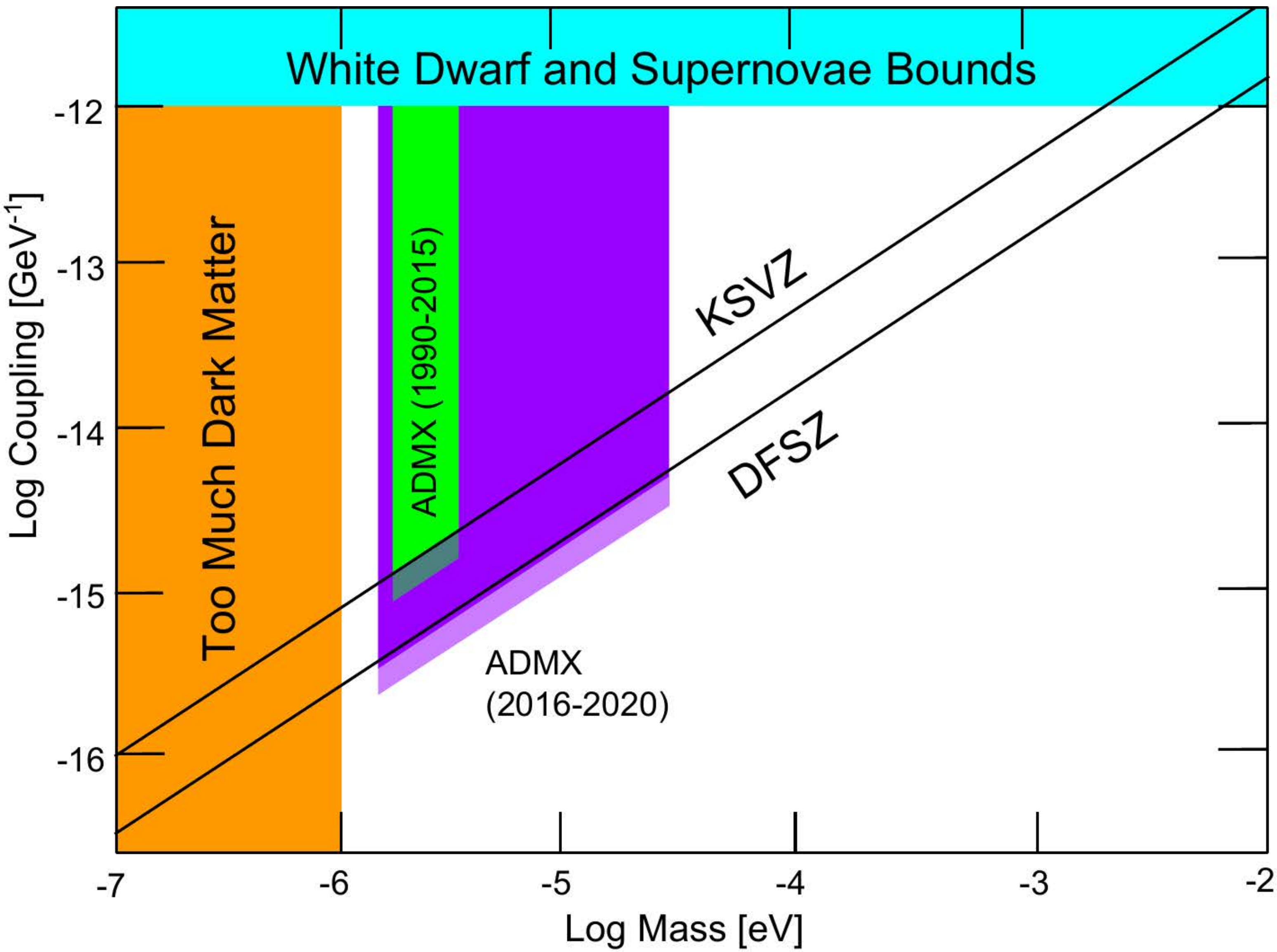
Tremmel et al 2017

Milky Way DM Distribution Function



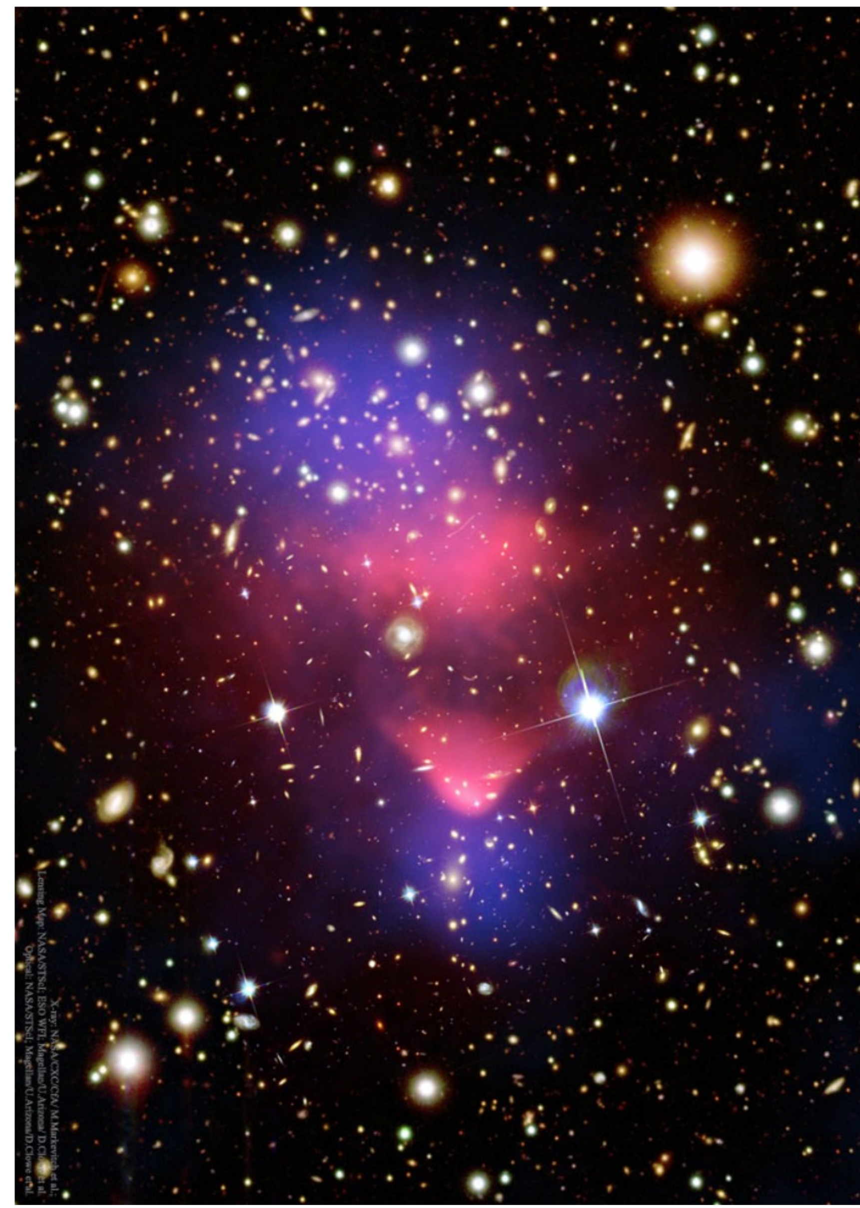
Erik Lentz

Consequences for DM Searches



Clusters: the science

- Largest bound objects in the Universe
- Visible across the entire Universe
- Baryonic content is observable
- “Closed box” for galactic evolution



Clusters: the challenge

- Good models of stellar feedback
- Good models of AGN (black hole) feedback
- Hydrodynamic instabilities require good algorithms
- Resolution: 10^5 Msun particles in 10^{15} Msun object
- Highly “clustered” computation

Charm Nbody GrAvity solver

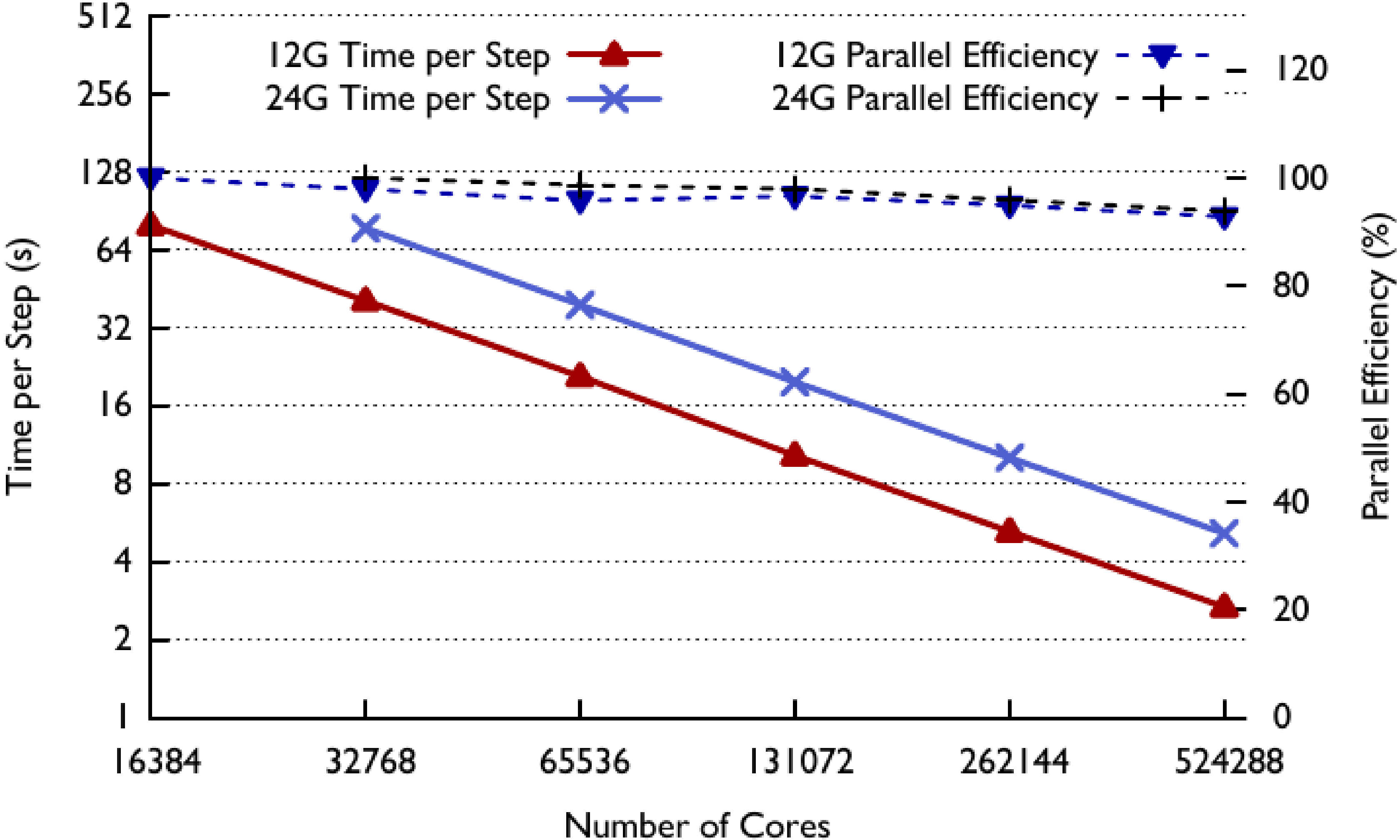
- Massively parallel SPH
- SNe feedback creating realistic outflows
- SF linked to shielded gas
- SMBHs
- Optimized SF parameters

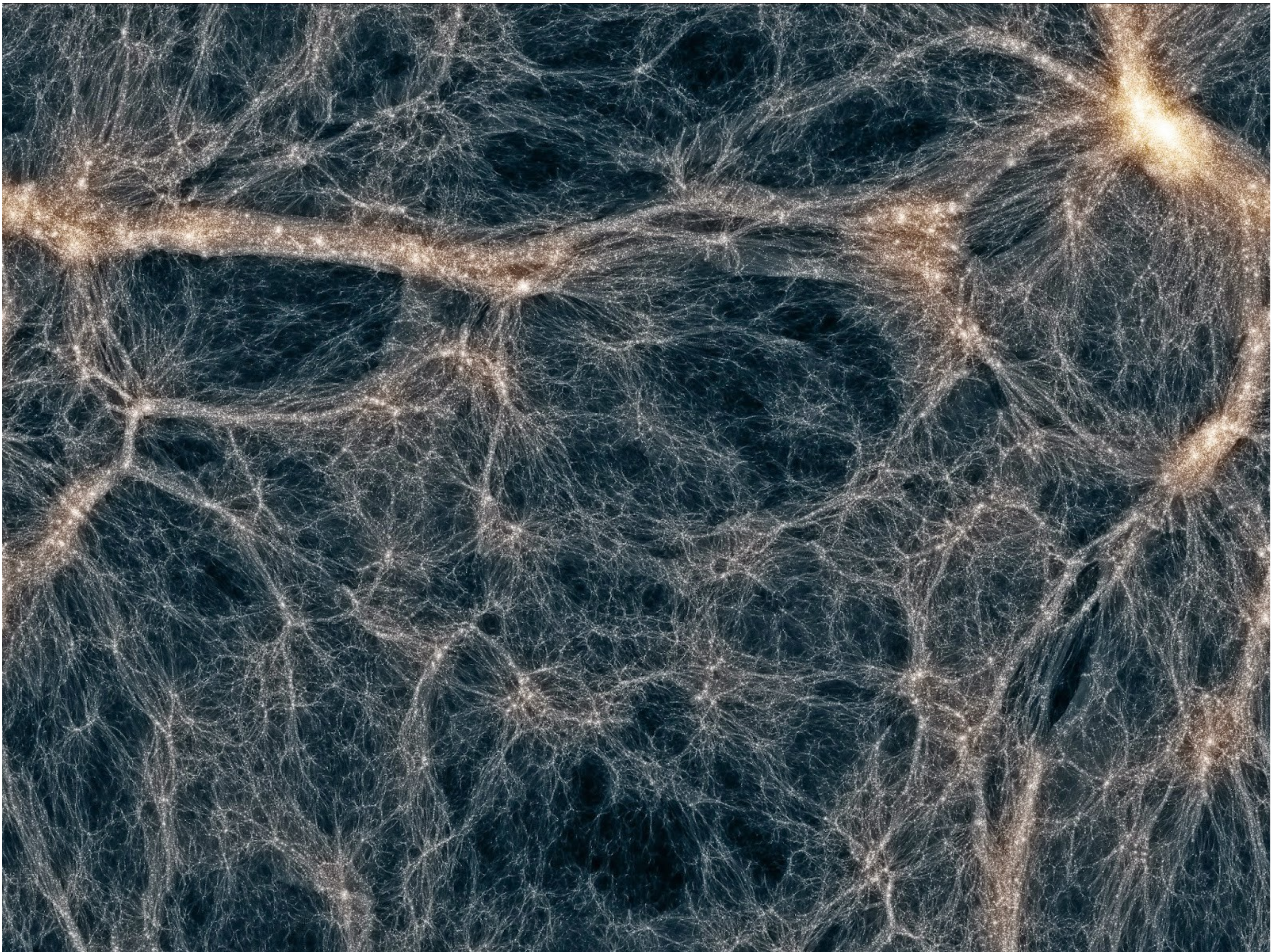
Menon+ 2014, Governato+ 2014

Charm++

- C++-based parallel runtime system
  - Composed of a set of globally-visible parallel objects that interact
  - The objects interact by asynchronously invoking methods on each other
- Charm++ runtime
  - Manages the parallel objects and (re)maps them to processes
  - Provides scheduling, load balancing, and a host of other features, requiring little user intervention

Scaling to .5M cores

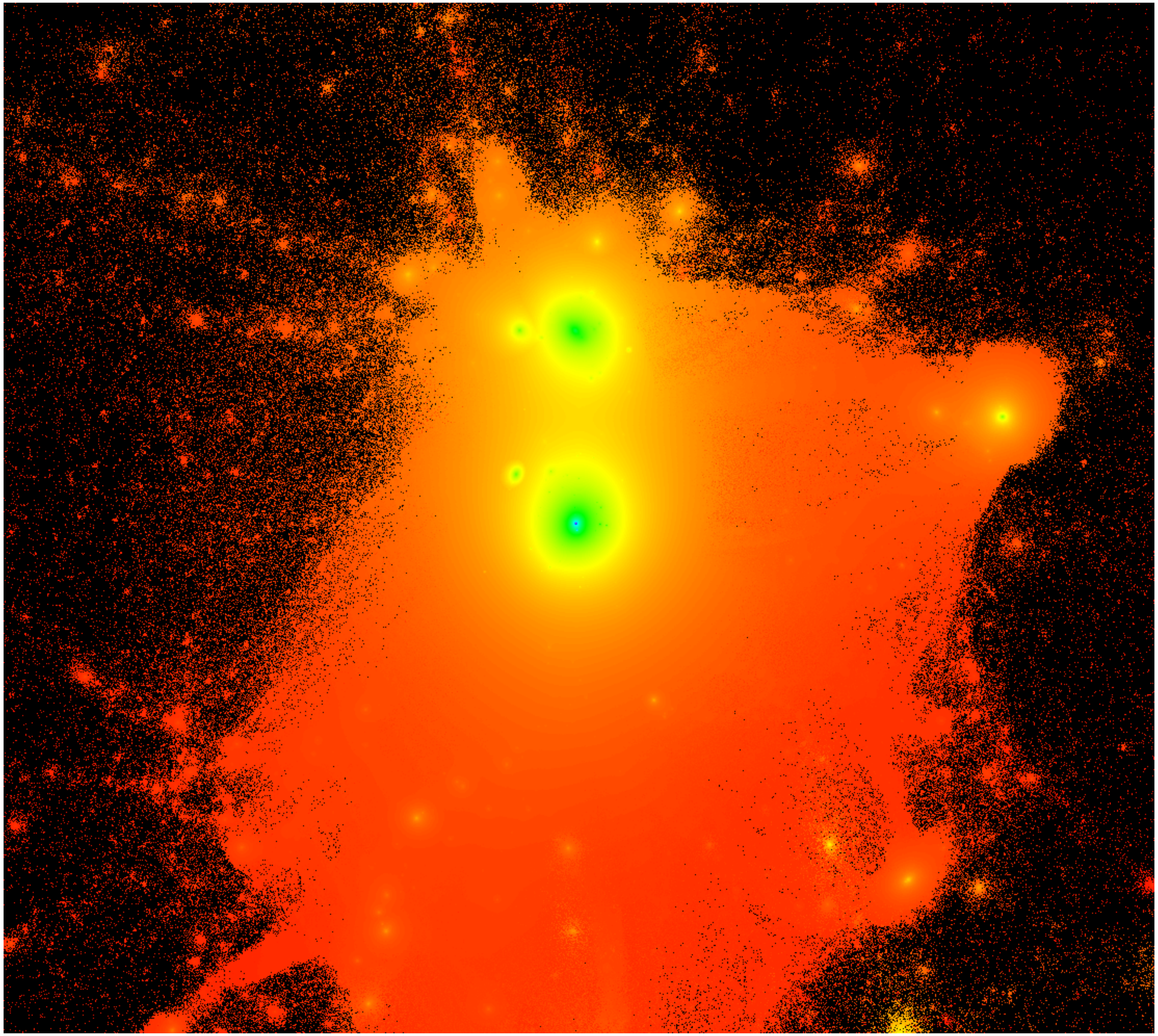




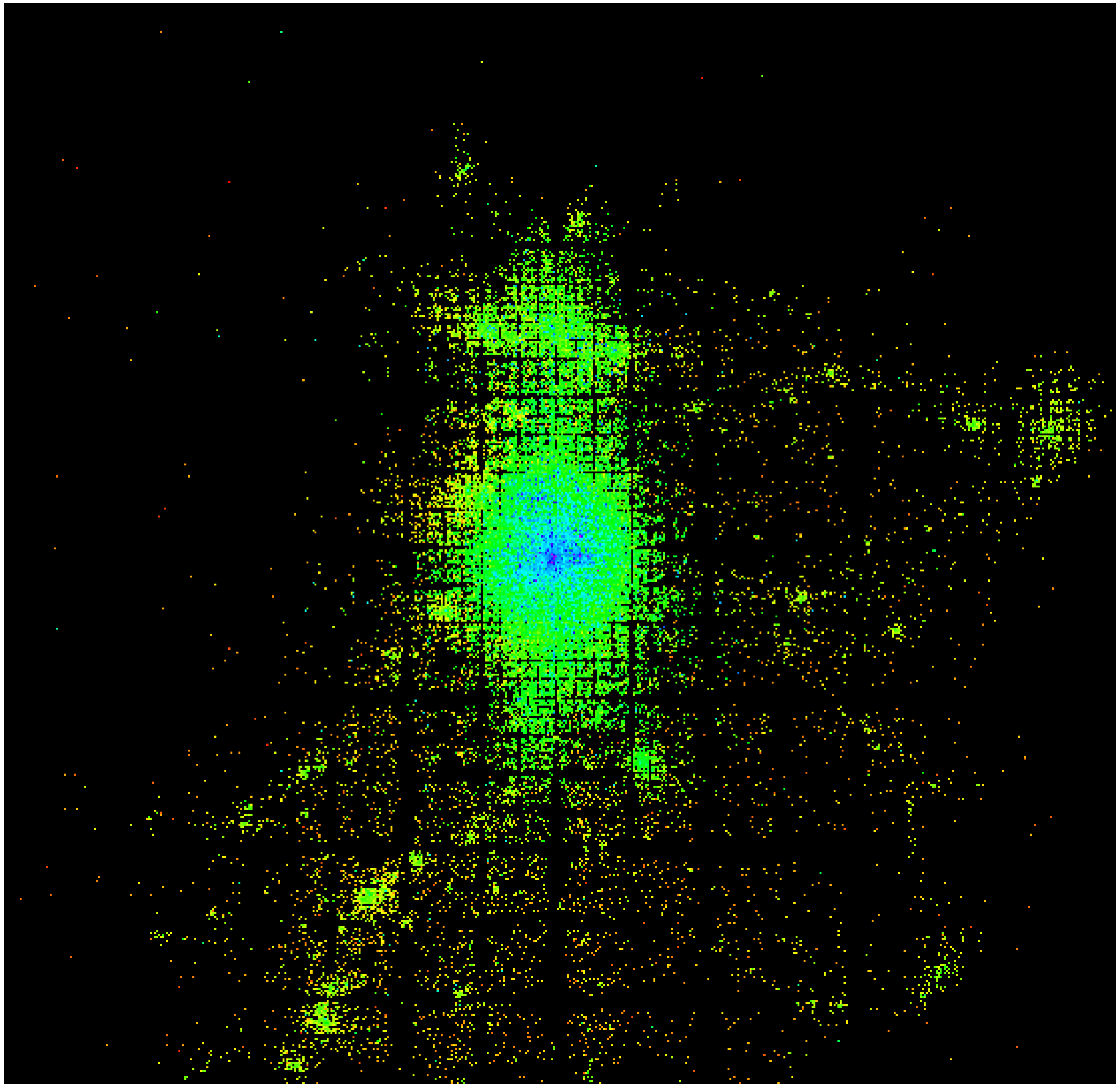
Clustered/Multistepping Challenges

- Computation is concentrated in a small fraction of the domain
- Load/particle imbalance
- Communication imbalance
- Fixed costs:
  - Domain Decomposition
  - Load balancing
  - Tree build

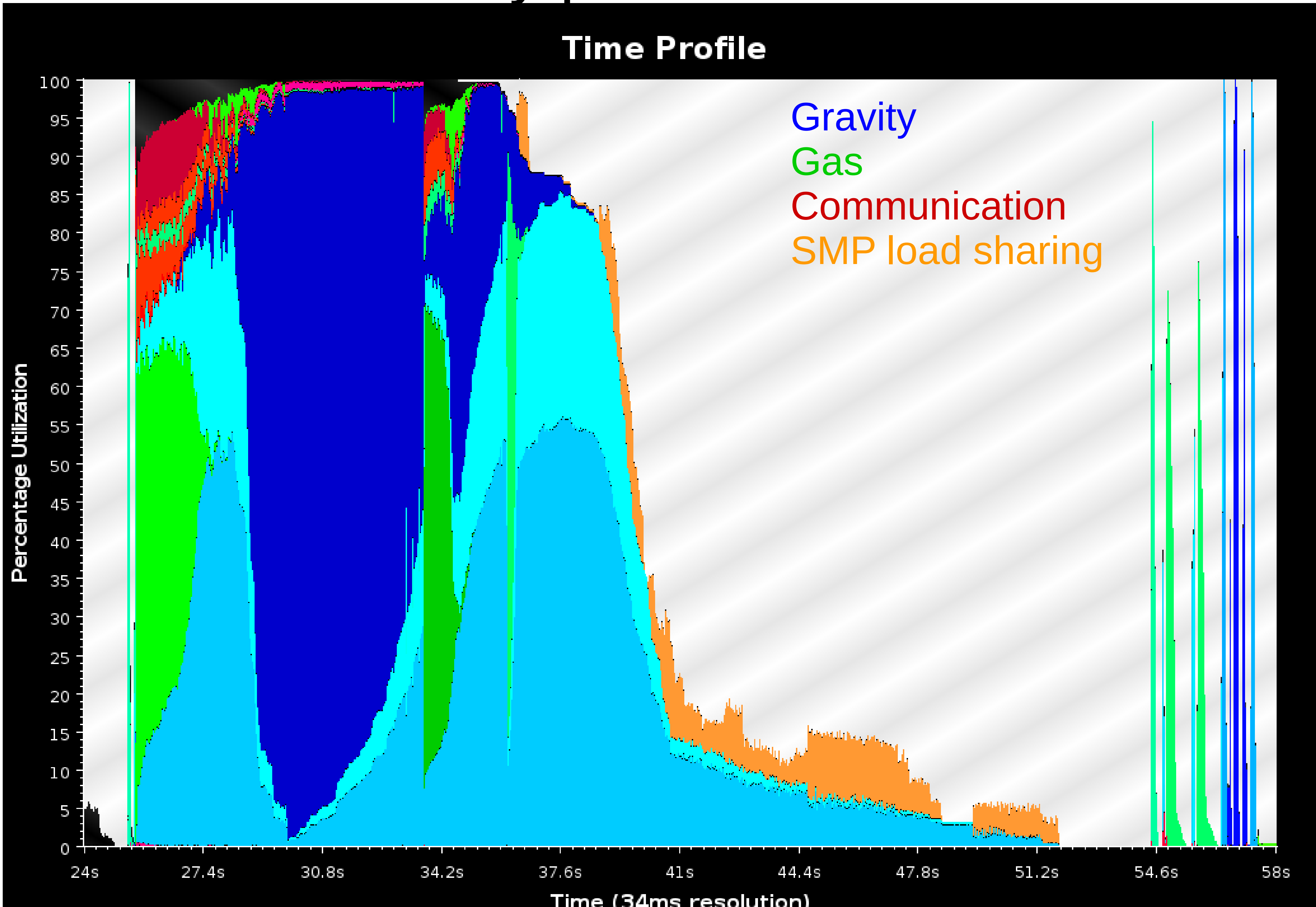
Zoomed Cluster simulation



Load distribution

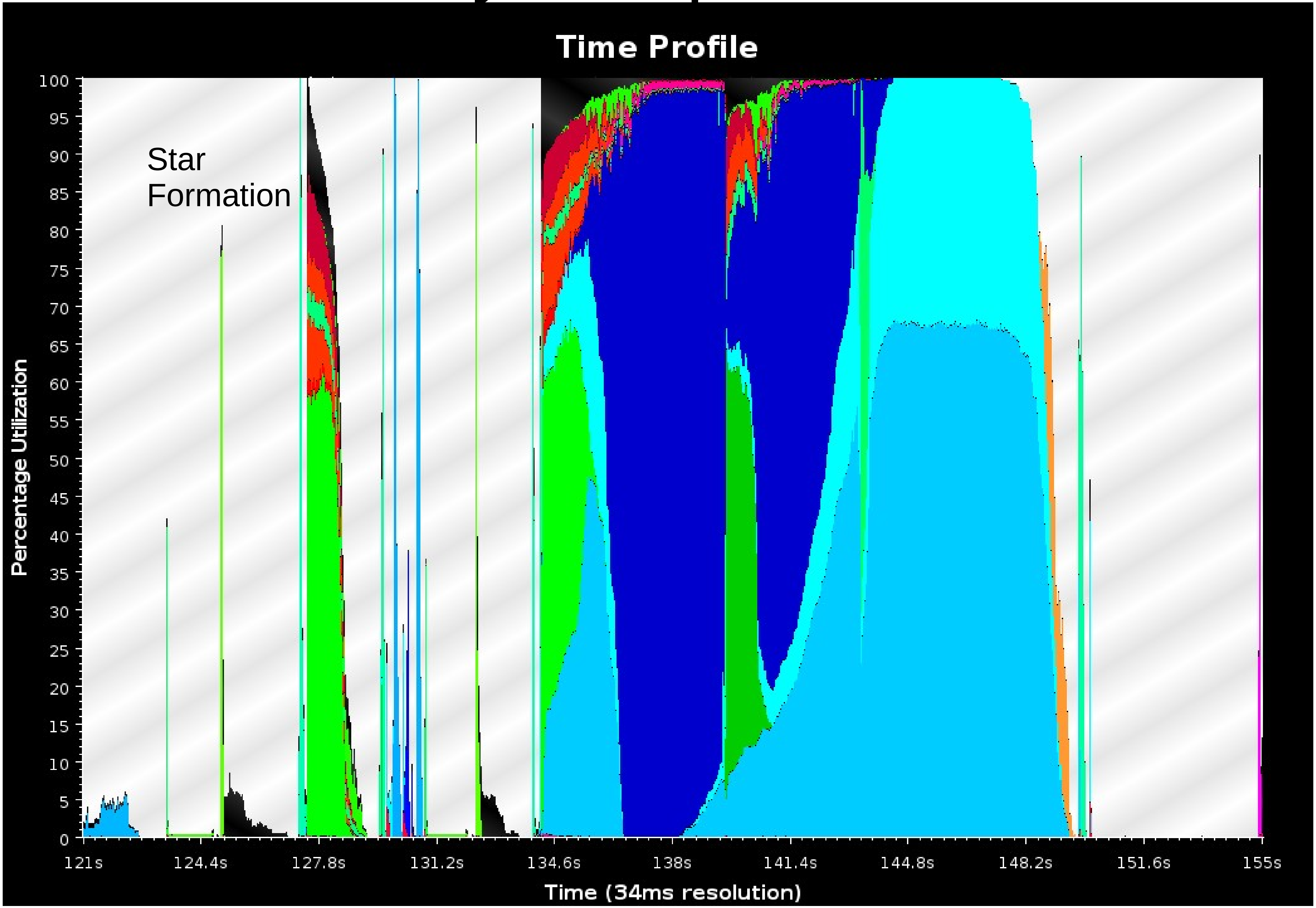


LB by particle count



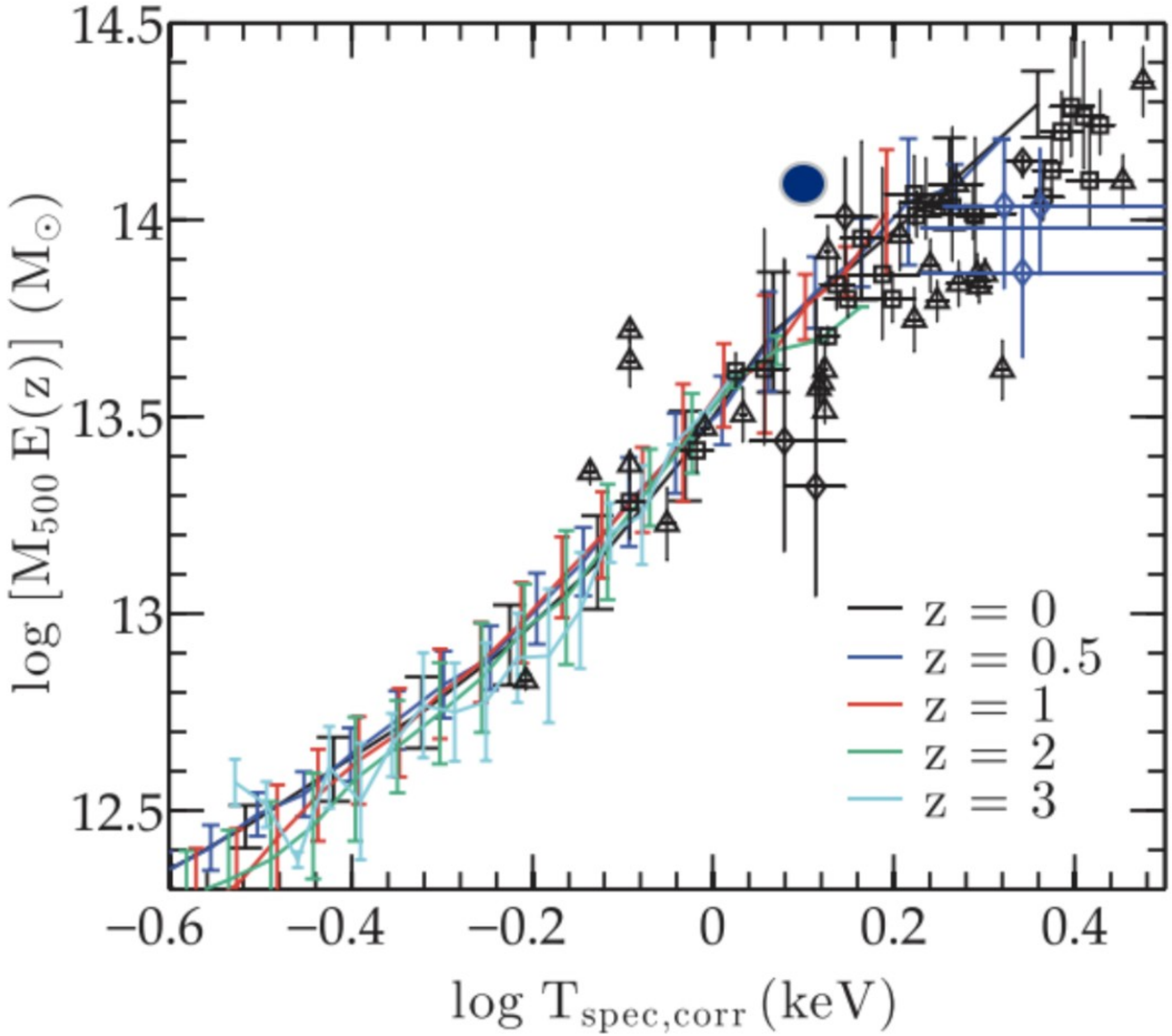
29.4 seconds

LB by Compute time



15.8 seconds

Preliminary Results: ICM Temperature



Take Aways

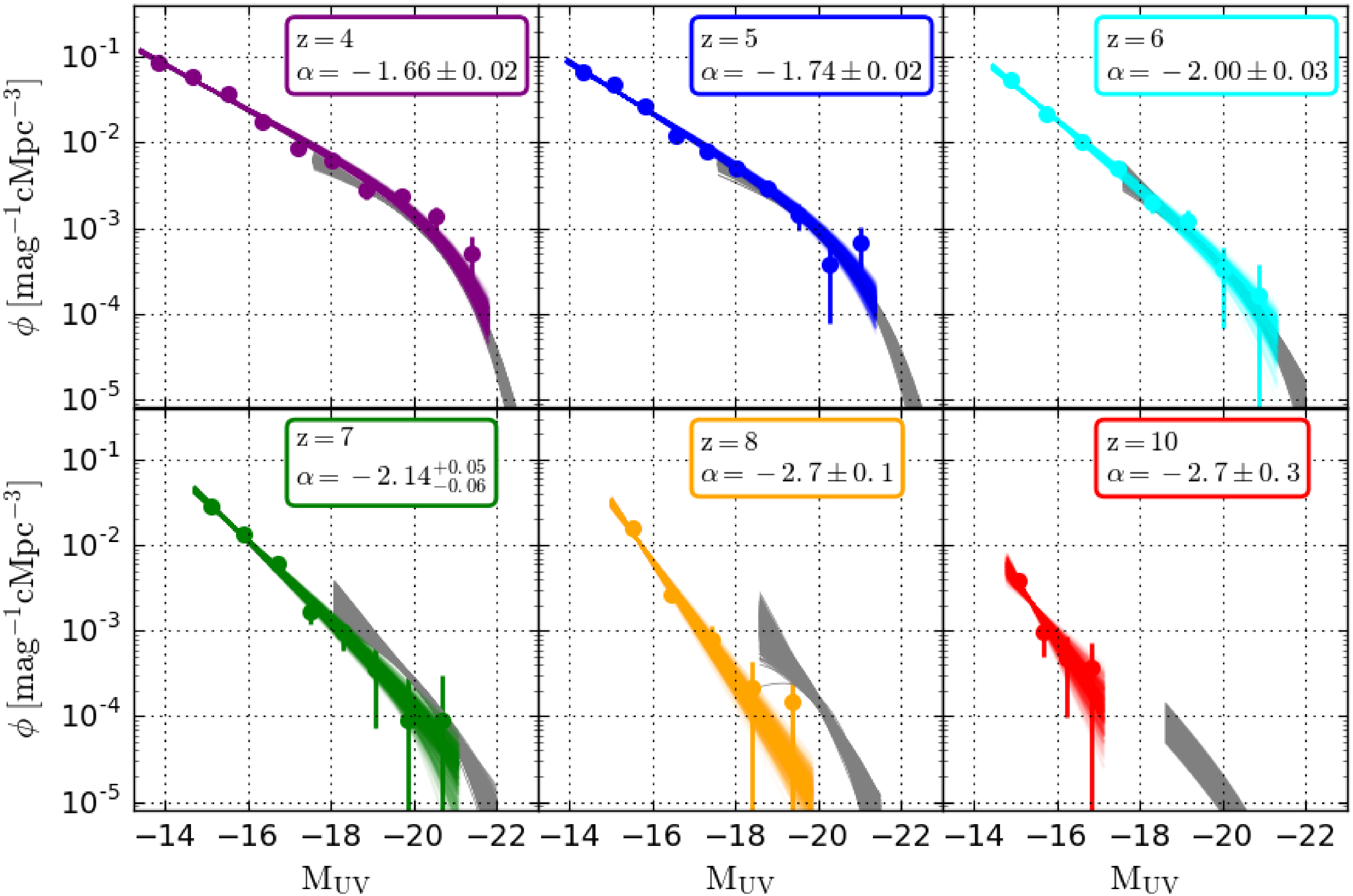
- Galaxy Clusters are hard:
  - Scale is set by galactic (i.e. star formation) physics
  - Orders of magnitude larger than galaxies
  - Computational effort is spatially concentrated.
  - (Probably should include magnetic fields; see **Iryna Butsky**'s talk)
- But clusters are doable
  - Capability machines
  - Advanced load balancing techniques

Acknowledgments

- NSF ITR
- NSF Astronomy
- NSF SSI
- NSF XSEDE program for computing
- BlueWaters Petascale Computing
- Blue Waters PAID Program
- NASA HST
- NASA Advanced Supercomputing

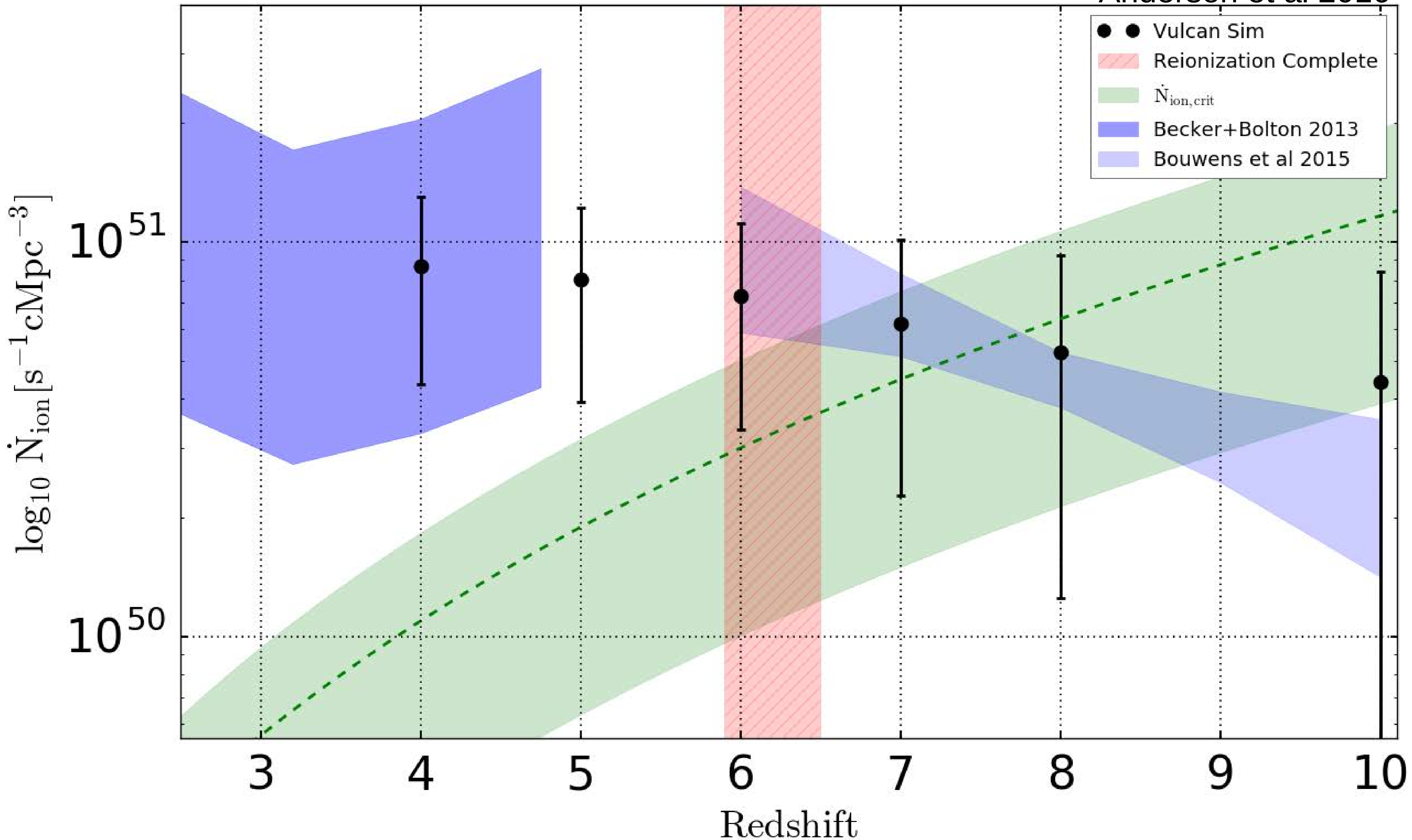
Luminosity Function

Anderson, et al 2016



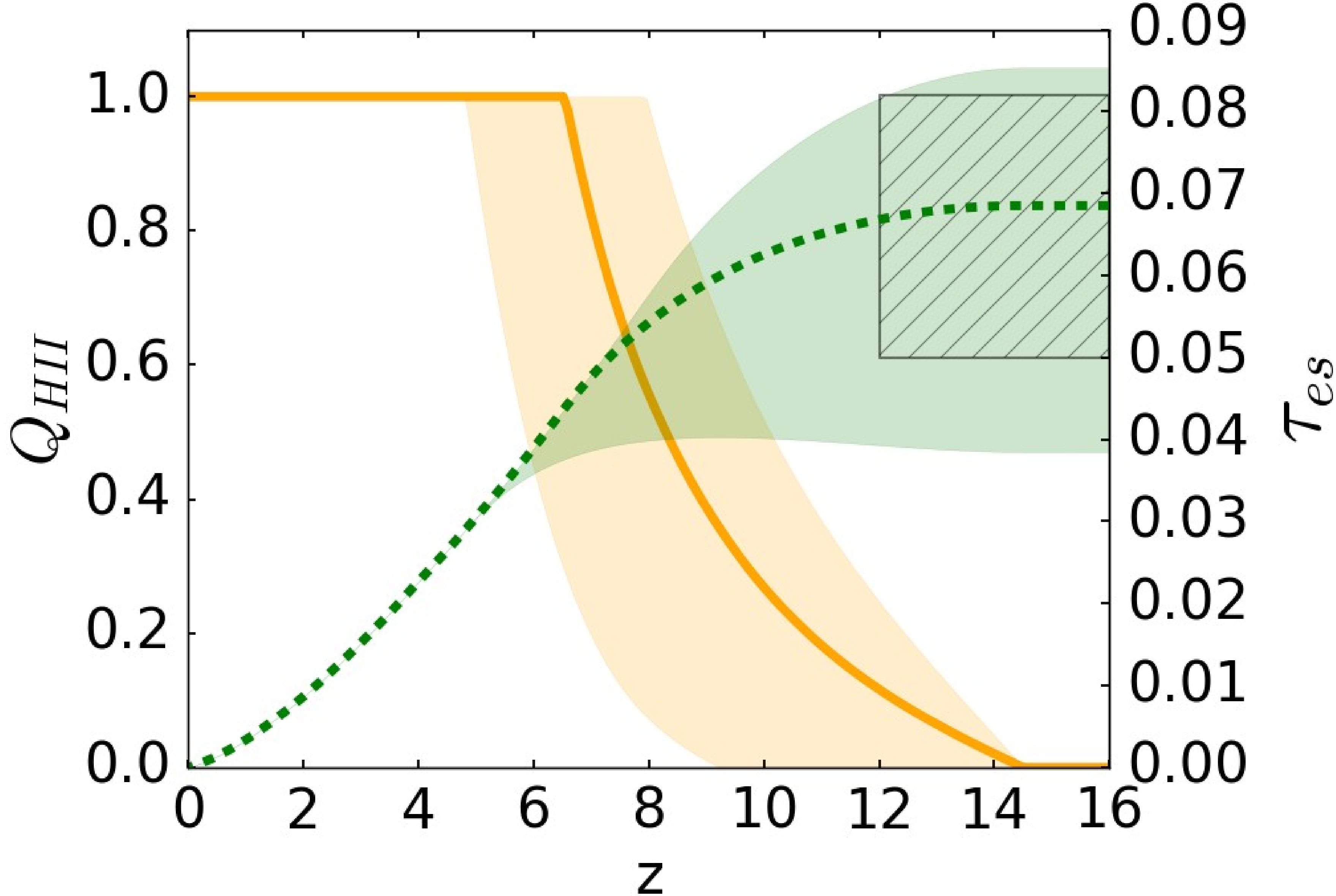
Faint galaxies reionize the Universe

Anderson et al 2016



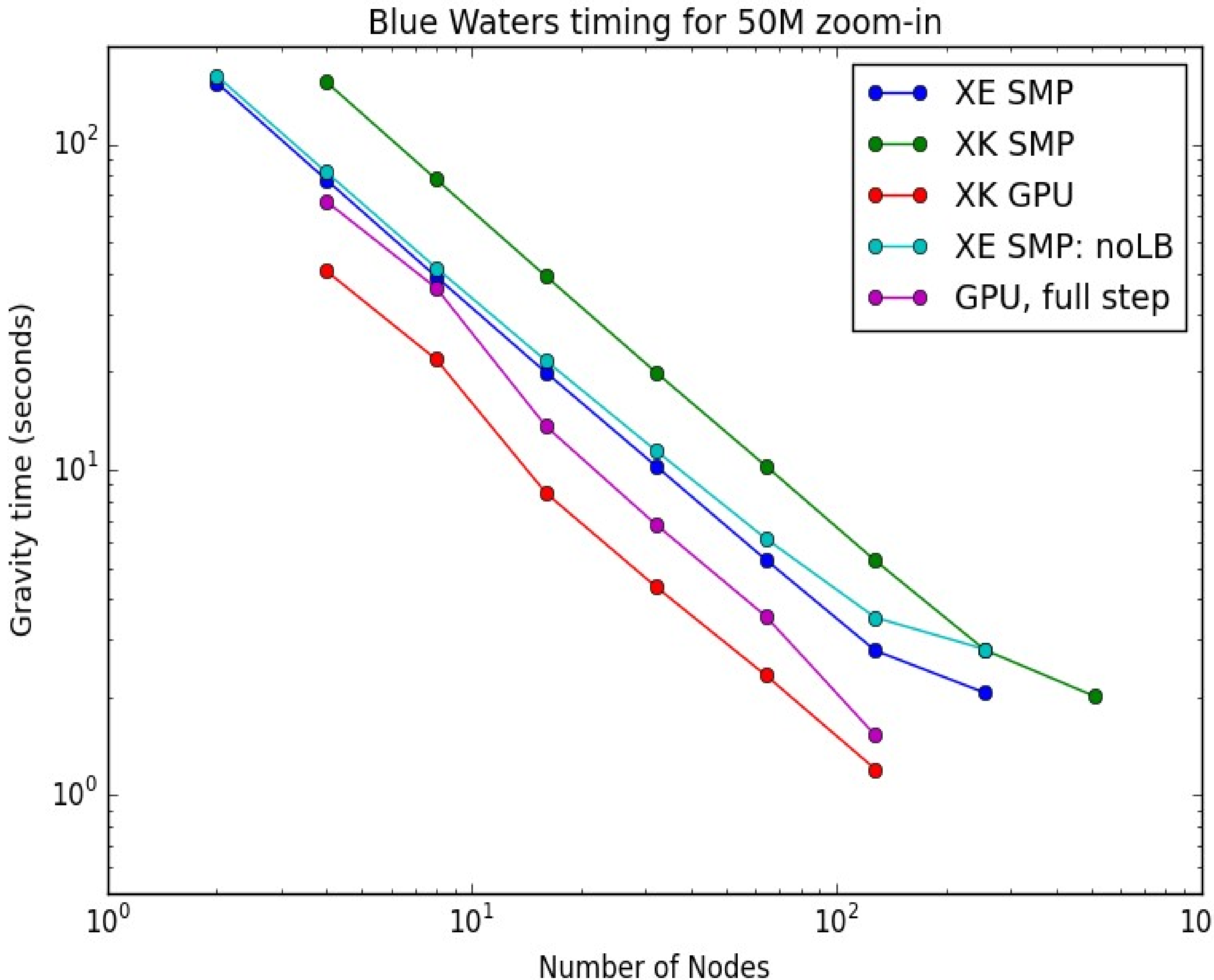
Faint galaxies reionize the Universe

Anderson et al 2016



PAID: ChaNGa GPU Scaling

- ChaNGa has a prelimary GPU implementation
- Goals of PAID:
  - Tesla → Kepler optimization
  - SMP optimization
  - Multistep Optimization
  - Load balancing
- Personnel:
  - Simon Garcia de Gonzalo, NCSA
  - Michael Robson, Harshitha Menon, PPL UIUC
  - Peng Wang, Tom Gibbs (NVIDIA)



Blue Waters timing for 50M zoom-in
XE SMP
XK SMP
XK GPU
XE SMP: noLB
GPU, full step
Gravity time (seconds)
Number of Nodes
10^0
10^1
10^2
10^3

PAID GPU Progress

- 2X speed up of main gravity kernel; 1.4X speedup of 2nd gravity kernel
  - Interwarp communication
  - Caching of multipole data
  - Higher GPU occupancy
  - Overall speedup of 60%
- SMP queuing of GPU requests
  - Reduced memory use, allowing more host threads
  - GPU memory management still an issue

Broader Impacts: Pre-Majors and Supercomputing

- UW Pre-Major in Astronomy Program:
  - Engage underrepresented populations in research early
  - Establish a cohort
  - Plug major leak in the STEM education pipeline
- Simulation data analysis is ideal for this research
  - Science and images are compelling
  - Similarity to Astronomical data reduction

Simulated Galaxy Catalogs

Zoe Deford
Joshua Smith
(UW Freshman)